

## A SPIN-MODULATED TELESCOPE TO MAKE TWO-DIMENSIONAL COSMIC MICROWAVE BACKGROUND MAPS

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### ABSTRACT

We describe the HEMT Advanced Cosmic Microwave Explorer (HACME), a balloon-borne experiment designed to measure subdegree-scale cosmic microwave background anisotropy over hundreds of deg<sup>2</sup>, using a unique two-dimensional scanning strategy. A spinning flat mirror that is canted relative to its spin axis modulates the direction of beam response in a nearly elliptical path on the sky. The experiment was successfully flown in 1996 February, achieving near laboratory performance for several hours at float altitude. A map free of instrumental systematic effects is produced for a 3.5 hr observation of 630 deg<sup>2</sup>, resulting in a flat-band power upper limit of  $\langle l(l+1)C_l/(2\pi) \rangle^{0.5} < 77 \mu\text{K}$  at  $l = 38_{-20}^{+25}$  (95% confidence). The experiment design, flight operations, and data, including atmospheric effects and noise performance, are discussed.

*Subject headings:* cosmic microwave background — cosmology: observations — instrumentation: detectors

### 1. INTRODUCTION

The cosmic microwave background (CMB) is a unique relic of the primordial universe, and the study of its properties promises to advance our understanding of the formation of structure in the universe. Detection of CMB anisotropy helps to discriminate among potential models with different cosmological parameters (Bond, Efstathiou, & Tegmark 1997; Jungman et al. 1996; Hu, Sugiyama, & Silk 1997; Zaldarriaga, Spergel, & Seljak 1997). Much of the discriminating power of CMB measurements comes from angular scales below 1°, or Legendre multipoles,  $l$ , greater than 100 in the angular power spectrum.

The effort to understand the power spectrum of the CMB for large  $l$  has continued since the discovery of anisotropy in the CMB (Smoot et al. 1992; Alsop et al. 1992; Schuster et al. 1993; Readhead et al. 1989; Gundersen et al. 1995; Lim et al. 1996; Cheng et al. 1997; Netterfield et al. 1997). Many previous experiments minimized errors for the best sensitivity at a single angular scale on a small region of sky. An ideal experiment will have sufficient sky coverage with enough sensitivity over a broad angular range to measure the power spectrum at several angular scales. Broadly speaking, the optimal strategy is to aim for a signal-to-noise ratio of order unity per pixel with the largest and most symmetric sky coverage possible, for a given resolution and instrument sensitivity (Knox 1997; Tegmark 1997a, 1997b). Instrumentation has evolved so that sensitive large-area maps can be obtained with subdegree resolution over a range of microwave frequencies (Coble et al. 1999; de

Oliveira-Costa et al. 1998; Netterfield et al. 1997). With sensitivity from the beam size to the large angle of the spinning chop, a multipixel spin-modulated telescope easily covers enough sky area to make precision measurements of the angular power spectrum. In principle, future experiments will enable cosmological parameter extraction from such power-spectrum measurements.

The HACME experiment is the precursor to several spin-modulated experiments that will accurately measure the angular power spectrum using cryogenic HEMT (high electron mobility transistor) amplifier receivers flown on balloon platforms to minimize atmospheric contamination and enable large-angle chopping. HACME has a single  $Q$ -band receiver fed by the ACME telescope (Meinhold et al. 1993) and modulated by a 1.2 m rotating canted flat mirror. The single horn and conventional duration balloon flight allows an upper limit on CMB anisotropy, and upcoming instruments will have more horns and longer flights to make power spectrum estimates. The novel optical spinning technique, described in detail in § 2.1, allows rapid two-dimensional modulation of the instrument to reduce the effects of  $1/f$  noise in the instrument while tying together widely separated regions of the sky in a short time. Changes in optics, computers, and data acquisition made for HACME are also described in § 2. In § 3, the observations made and the data collected are discussed.

### 2. INSTRUMENT

#### 2.1. Optics and Scan Strategy

HACME's optical system (Fig. 1) consists of a fixed ellipsoidal mirror, the original ACME 1 m parabolic primary

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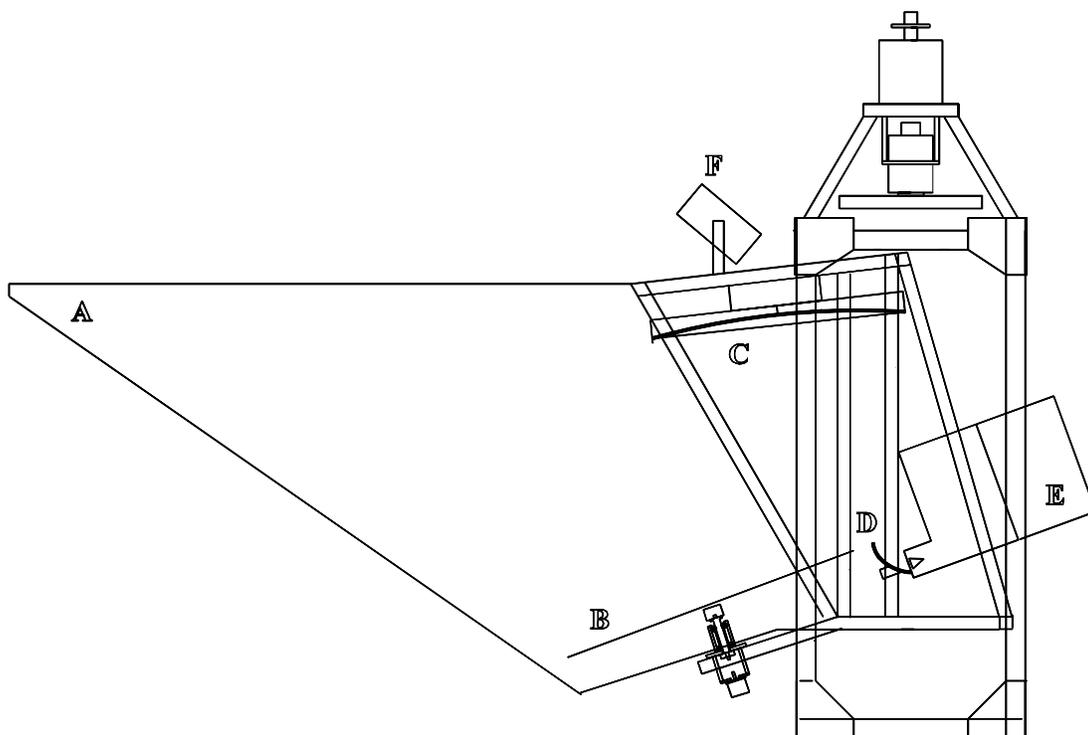


FIG. 1.—Schematic drawing of the major optical elements of HACME: A: Baffle; B: Canted spinning flat mirror; C: 1 m parabolic mirror; D: Ellipsoidal mirror; E: Dewar; F: CCD camera.

mirror, a new 1.2 m spinning flat mirror, and a baffle surrounding the flat mirror between  $2^\circ$  and  $5^\circ$  from the optical axes to intercept direct spillover paths from the Earth to the edge of the primary mirror. The spinning flat mirror, canted  $2.5^\circ$  relative to its spin axis, modulates the telescope beam in a nearly elliptical path on the sky, so we become sensitive to anisotropy on scales from the beam size,  $0.8$  FWHM, to the size of the major axis of the ellipse,  $10.0$ . The polar angle,  $\theta_r$ , and an azimuthal angle about the center of the path,  $\phi_r$ , are given by

$$\cos(\theta_r) = \cos(2\theta_c) + 2 \sin^2(\theta_r) \sin^2(\theta_c) \sin^2(\phi_f), \quad (1)$$

$$\tan(\phi_r) = \sin(\phi_f) \times$$

$$\left[ \frac{\cos(\theta_c) \cos(\theta_r) - \sin(\theta_c) \sin(\theta_r) \cos(\phi_f)}{\cos(\theta_c) \cos(\phi_f) + (1/2) \sin(2\theta_r) \sin(\theta_c) \sin^2(\phi_f)} \right], \quad (2)$$

where  $\theta_c = 2.5^\circ$  is the cant,  $\theta_r = 31.9^\circ$  is the angle between the mirror spin axis and the main optical axis, and  $\phi_f$  is the phase of the flat-mirror spin. Performing an azimuth scan adds another level of modulation, building up a two-dimensional image of the CMB. At the start of the azimuth scan, the leftmost part of the beam path is aligned with a fiducial star. The gondola scans  $20^\circ$  to the right and returns at a rate of  $0.17 \text{ s}^{-1}$ . This process continues as we track the fiducial star across the sky. This observing strategy is shown on a map in Figure 2. Data taken in this manner produce large maps with nearly diagonal noise covariance matrices (Wright 1996; Tegmark 1997b). The resulting region of sky covered is almost  $30^\circ$  wide and between  $10^\circ$  and  $30^\circ$  high for a typical 3–5 hr observation.

Telescope response is measured on an antenna range in Santa Barbara with a Gunn diode oscillator source. The best-fit Gaussian to the main lobe along the azimuthal and

elevation axes at 40 GHz are  $0.84 \pm 0.02$  and  $0.75 \pm 0.02$  FWHM, respectively. The beam shape is independent of the orientation of the flat mirror to within the measurement errors and varies as the inverse of the frequency of the subbands described below. Observations of the Moon during flight confirm our ground tests of the beam size. We

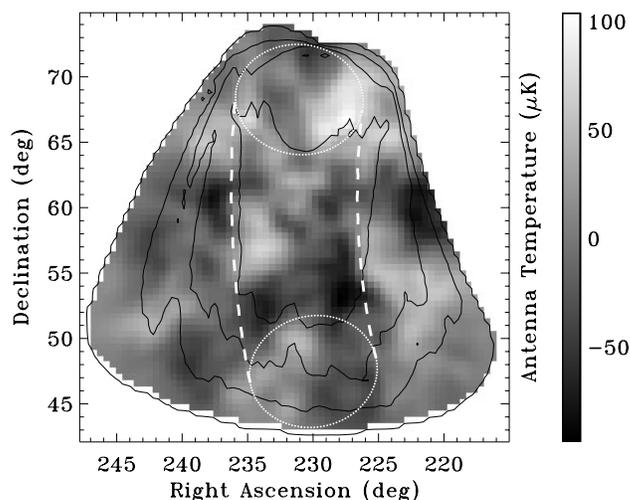


FIG. 2.—Map of  $\gamma$  UMi. The map, made from the 43 GHz subband data, is consistent with instrumental noise. There are over 3000 gnomonic  $28'$  square pixels. The contours encompass regions covered greater than 12.8, 6.4, and 3.2 s of integration per  $\text{deg}^2$ . The white dots represent the 125 samples taken along an elliptical path for each spin of the flat mirror. Spins at each end of the azimuth scan are shown along with the dashed boundary of the region observed during one scan of the telescope.

measured the illumination taper of the primary to be below  $-25$  dB around the edge and calculated the taper around the rotating flat mirror to be always less than  $-39$  dB. Sidelobe measurements made with the main baffle in place demonstrate response below  $-60$  dB for angles greater than  $20^\circ$  from boresight along the elevation axis.

### 2.2. Receiver and Data Acquisition

The HACME  $Q$ -band receiver is the same as that of SP94 (Gundersen et al. 1995), with a five-stage InP-GaAs HEMT multiplexed into three 2.3 GHz bandwidth subbands centered at 39.15, 41.45, and 43.75 GHz, but is used as a total-power radiometer for this instrument. The output diodes for each subband connect to a low-noise preamplifier, an A/C coupling stage ( $1/RC \approx 7$  s), and a linear voltage to frequency (V/F) converter. The output of the V/F is counted synchronously with the spin of the flat mirror, 125 triggers per rotation. An indexed optical encoder attached to the spinning mirror shaft provides the triggers that divide each mirror spin into 125 “sectors.” The mirror spins at  $2.500 \pm 0.004$  Hz, resulting in an ideal integration of the diode outputs at 3.2 ms. Pulses from a 16 MHz clock are counted for normalization. Pointing and receiver data are stored on board the payload, and a compressed form of the data is sent to the ground. Analysis is performed on the data recovered from the hard disk after landing.

### 2.3. Calibration

The telescope is calibrated using two methods. First, the absolute calibration, including atmospheric attenuation and telescope loss, is determined by observing the Moon. A model of lunar emission by S. Keihm (JPL) is compared to two-dimensional maps made from the data. Second, an ambient-temperature “Eccosorb” target can intercept the beam between the primary and secondary mirrors to provide relative calibration during the flight. The latter calibrations are performed about every hour during flight and determine the relative calibration to  $\pm 3\%$ . The following results are calibrated in antenna temperature units. Conversion of the data to CMB thermodynamic temperature requires multiplication by 1.04, 1.05, and 1.05 for the 39, 41, and 43 GHz subbands, respectively.

## 3. DATA

### 3.1. Observations

This experiment flew for the first time on 1996 February 11 from Fort Sumner NM, obtaining 15 hr at a float altitude of 35.4 km. Operations proceeded smoothly, with detectors and all servo systems reaching expected performance for the entire flight. Of the 12 nighttime hr at float, 8 hr were spent mapping three extended regions near the stars  $\alpha$  Leo,  $\gamma$  Eri, and  $\gamma$  UMi. These regions were chosen for low foreground emission and with systematic tests in mind. All regions observed are between  $40^\circ$  and  $50^\circ$  from the Galactic plane. When observed,  $\alpha$  Leo was rising in the east,  $\gamma$  Eri was setting in the southwest, and  $\gamma$  UMi was transiting in the north. The remainder of the time was spent on calibrations, main lobe mapping using the Moon, and systematic tests such as operating without the servo systems and RF transmitters. A second flight was attempted on 1996 June 1 but was terminated by a balloon failure before reaching float altitude.

### 3.2. Flight Performance

During flight, the HACME telescope performed as expected with no known difficulties. Once scanning a region, the telescope collects data very efficiently. For the  $\gamma$  UMi observation, after removal of 16% of the data due to calibrations, we are able to use 99% of the remainder. One of the advantages of a balloon environment is that essentially all of the data is usable, since there is no “bad weather.” The power spectra of noise for the three subbands have high-frequency (white-noise) limits of 0.9, 0.5, and 0.4  $\text{mK s}^{1/2}$ , with a typical knee frequency of 60 Hz (the frequency at which the noise power is double the white-noise limit). The subbands have typical correlations between 0.5 and 0.7 due to HEMT amplifier gain fluctuations. This performance is consistent with our laboratory measurements of this receiver and is typical of cryogenic HEMT amplifiers.

### 3.3. Systematics

#### 3.3.1. Synchronous Offset

In order to make maps from the data, we need to remove systematic offsets. When the data are binned into “sectors” rather than sky pixels, we find an anomalous spin-synchronous offset for each subband. The peak-to-peak amplitudes of the offsets are 1.3, 1.6, and 4.0 mK in the 39, 41, and 43 GHz subbands, respectively. The amplitude of these signals varies less than  $5\% \text{ hr}^{-1}$  over the entire flight. Among the possible causes are atmospheric emission, thermal emission from the optics, the CMB dipole, and far-sidelobe contamination. Because the observed offset varies slowly, we are able to remove it on short timescales. We discuss the possible contributions to this offset and its removal below.

While the atmosphere is the dominant cause of a spin-synchronous offset on the ground, we find that it is small at our float altitude and remarkably consistent with our predictions at lower altitudes based on a model for atmospheric emission. We verify the spectrum of the atmospheric contribution and its elevation and altitude dependence up to 25 km with data taken during the ascent of the balloon payload (Fig. 3). At an altitude of 35 km, the largest peak-to-peak signals predicted are 500, 600, and 800  $\mu\text{K}$  for our detection subbands. During the  $\alpha$  Leo measurement, the telescope tracked a  $22^\circ$  range in elevation, over which the atmospheric signal would have doubled. With no detectable monotonic change in the offset, the atmospheric signal in the 43 GHz band is limited to less than 600  $\mu\text{K}$  for this observation.

Thermal emission from the optics is expected to be a larger contribution to the spin-synchronous offset and opposite in phase from the atmosphere. Our  $Q$ -band receiver detects only one linear polarization, and a synchronous signal is expected as the angle of the flat mirror relative to the receiver changes (Wollack et al. 1997). For our horizontal  $E$ -plane, the signal is largest when the beam elevation is greatest. This signal depends on the temperature and emissivity of the flat mirror. The peak-to-peak amplitude of this signal is predicted to be 7.4, 7.6, and 7.8 mK in the increasing frequency subbands. The polarization signal is calculated to decrease just 10% over the course of the entire flight as the flat mirror cools.

We calculate the signal expected from the motion-induced Doppler dipole anisotropy (Lineweaver et al. 1996). For a  $10^\circ$  throw, the maximum possible signal from the

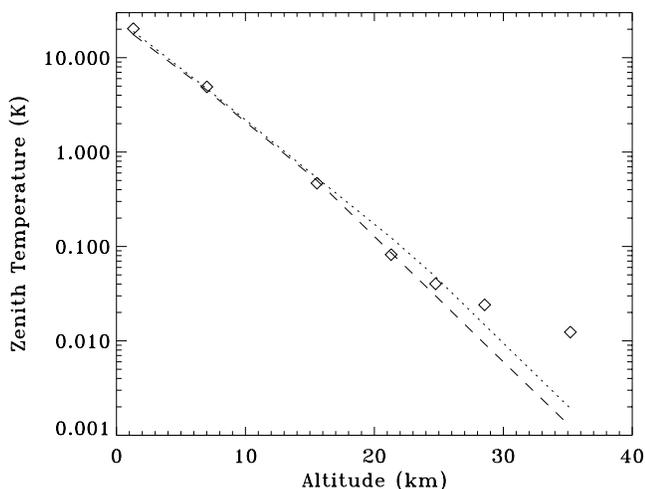


FIG. 3.—Spin-synchronous offset in 43 GHz subband vs. altitude. Diamonds show the zenith temperatures assuming that the detected offset is due only to the atmosphere. The dotted line shows an atmospheric model prediction using Zhevakin-Naumov line shapes. The dashed lines show Van Vleck-Weisskopf line shapes. The observation and models differ at float altitude, because the atmosphere is no longer the dominant contribution to the offset.

dipole measured by HACME would be  $580 \mu\text{K}$ . The spin-synchronous offset produced by the dipole is equal in the three subbands, and its phase and magnitude vary with the equatorial coordinates of the target.

A linear combination of the sources above cannot explain the observed offsets throughout the flight, however, so other effects must be present. Far-sidelobe contamination, radio-frequency interference, chopped spillover, and electrical interference are possibilities. Because the Earth makes the largest contribution to far-sidelobe response, we expect this systematic error to be elevation dependent. Because the spin-synchronous offset did not change appreciably during the  $\alpha$  Leo observation, sidelobe contamination is not likely to be a major cause of systematic error. Radio-frequency interference was of great concern to us, due to the necessity of the NSBF flight transmitters and our previous experience. We were, however, able to turn off and on all on-board transmitters (we even obtained several minutes with *all* transmitters off), and saw no effect on the data. We cannot entirely rule out chopped spillover and electrical pickup. Chopped spillover refers to near-field off-axis telescope response modulated by the flat mirror or its balance arm. We performed careful noise integration tests prior to flight to levels far below the observed offset, but differences between ground tests and flight could lead to systematic errors from electrical interference.

### 3.3.2. Offset Removal

Current data do not allow the separation and identification of the modeled contributions partly because of the cancellation of the effects. A conservative approach to offset removal is taken because the offsets detected are due to a superposition of the contributions listed above. We remove all data at the first and second harmonic of the spin rate before analysis, since the offset is expected and observed to be significant only at these harmonics. Only 3.2% of the 2–3 million degrees of freedom in an observation are sacrificed

to this harmonic removal. We find it a robust approach that sacrifices sensitivity to multipoles  $l < 20$  and use it for all following analysis.

### 3.4. Maps

The total power data collected by HACME are easily converted into sky maps. Measurements of each “sector” are assigned a sky position using equations (1) and (2) and the azimuth and elevation of the center of each flat mirror spin. Averaging measurements of sky pixels results in maps dominated by stripes due to  $1/f$  noise and by the synchronous offset. The map-making process in Tegmark (1997a) and Wright (1996) is needed to reduce striping and excess noise due to the amplifier gain fluctuations, and is convenient for the offset removal. This algorithm accounts for the correlations introduced by  $1/f$  noise and optimally reduces striping in the maps by using the interconnections of our scan pattern. The same pointing information is used along with  $1/f$  noise estimates to make improved maps. Wiener filtering using signal-to-noise ratio eigenmodes can be used to project out the map modes with the best sensitivity.

Application of this method to these data is described in Tegmark et al. (1999, hereafter T99). Figure 2 shows a Wiener-filtered map from T99 of the  $630 \text{ deg}^2$ , 3.5 hr  $\gamma$  UMi observation, with integration time contours and the scan pattern overlaid. All structure seen in the map is consistent with instrument noise with pixel errors a factor of 2.6 greater than those expected for the white-noise levels indicated above. The sensitivity is best in the inner, more heavily sampled regions. T99 find a flat-band power upper limit of  $\langle l(l+1)C_l/(2\pi) \rangle_B^{0.5} < 77 \mu\text{K}$  (95% confidence) at  $l = 38_{-20}^{+25}$  from this observation.

## 4. CONCLUSION

We have built, tested, and flown a novel experiment to make large sky area two-dimensional maps of the CMB. The new equipment used at balloon altitudes, the HEMT amplifier and canted spinning flat mirror, worked flawlessly. We find that atmospheric emission should not significantly contaminate data of future experiments. The maps produced yield an upper limit to CMB anisotropy because of poorer sensitivity than expected. They show no systematic effects or striping from the  $1/f$  noise and demonstrate that the scan strategy used will be an excellent way to make more CMB anisotropy maps.

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