

SEARCH FOR UNRESOLVED SOURCES IN THE COBE-DMR TWO-YEAR SKY MAPS<sup>1</sup>A. KOGUT,<sup>2</sup> A. J. BANDAY,<sup>3</sup> C. L. BENNETT,<sup>4</sup> G. HINSHAW,<sup>2</sup> K. LOEWENSTEIN,<sup>2</sup> P. LUBIN,<sup>5</sup>  
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## ABSTRACT

We have searched the temperature maps from the COBE Differential Microwave Radiometers (DMR) first two years of data for evidence of unresolved sources. The high-latitude sky ( $|b| > 30^\circ$ ) contains no sources brighter than  $192 \mu\text{K}$  thermodynamic temperature ( $322 \text{ Jy}$  at  $53 \text{ GHz}$ ). The cumulative count of sources brighter than threshold  $T$ ,  $N(>T)$ , is consistent with a superposition of instrument noise plus scale-invariant spectrum of cosmic temperature fluctuations normalized to  $Q_{\text{rms-PS}} = 17 \mu\text{K}$ . We examine the temperature maps toward nearby clusters and find no evidence for any Sunyaev-Zel'dovich effect,  $\Delta y < 7.3 \times 10^{-6}$  (95% CL) averaged over the DMR beam. We examine the temperature maps near the brightest expected radio sources and detect no evidence of significant emission. The lack of bright unresolved sources in the DMR maps, taken with anisotropy measurements on smaller angular scales, places a weak constraint on the integral number density of any unresolved Planck-spectrum sources brighter than flux density  $S$ ,  $n(>S) < 2 \times 10^4 (S/1 \text{ Jy})^{-2} \text{ sr}^{-1}$ .

*Subject headings:* cosmic microwave background — radio continuum: general

## 1. INTRODUCTION

The COBE-DMR sky maps represent an unbiased survey of the full sky at frequencies 31.5, 53, and 90 GHz (wavelengths 9.5, 5.7, and 3.3 mm) and  $7^\circ$  angular resolution. Possible unresolved astrophysical sources at these wavelengths include Galactic H II regions, extragalactic radio and infrared sources, inverse Compton scattering of cosmic microwave background (CMB) photons by hot gas in clusters of galaxies (Sunyaev & Zel'dovich 1980), as well as CMB anisotropy from varying gravitational potentials on the surface of last scattering (Sachs & Wolfe 1967).

Bennett et al. (1992b, 1993) limit the total contribution from unresolved noncosmic sources averaged over the high-latitude ( $|b| > 30^\circ$ ) sky by comparing the first-year DMR maps to maps of known source distributions. With the exception of the Galactic quadrupole, there is no evidence for significant noncosmological contribution (e.g., not arising from temperature variations in the cosmic microwave background) to the anisotropy in the DMR sky maps. Since these results depend on an average over a large area of the sky, they do not preclude detection of weak noncosmological emission restricted to smaller regions of the sky. A pixel-by-pixel search for sources in the first year DMR data shows no evidence for unresolved sources at the  $200 \mu\text{K}$  level (Bennett et al. 1993). We report here

further results of a search for a population of weak unresolved sources in the two-year DMR maps, using both pixel-by-pixel techniques and a statistical comparison of the maps at positions where noncosmological emission is expected to be strongest.

## 2. FULL-SKY SEARCH

The DMR by design is insensitive to unresolved sources. The instrument response  $T_A$  to a source temperature distribution  $T(\theta, \phi)$  is

$$T_A = \frac{\int d\Omega T(\theta, \phi) P(\theta, \phi)}{\int d\Omega P(\theta, \phi)},$$

where we approximate the inner portion of the antenna pattern  $P(\theta, \phi)$  as a Gaussian with  $7^\circ$  full width at half-maximum. Sources with angular size smaller than the beam area (solid angle  $\sim 69$  square degrees; Bennett et al. 1992a) are diluted by the ratio of the source area to beam area. Flux densities 65, 180, and 536 Jy are required to produce a  $100 \mu\text{K}$  signal in the 31, 53, and 90 GHz channels, respectively.

The DMR maps are dominated by instrument noise and nearly scale-invariant CMB anisotropy (Smoot et al. 1992; Bennett et al. 1994). The DMR beam is wider than the  $2.6$  diameter pixels used to map the temperature anisotropy. Unresolved sources in the DMR maps can be distinguished from noise peaks by their correlated signal in neighboring pixels. We combine the two channels, A and B, at each DMR frequency to form “sum”  $(A+B)/2$  and “difference”  $(A-B)/2$  maps. For each pixel of these maps, we subtract the mean of all other pixels within  $10^\circ$  as a local baseline, then fit the patch to a Gaussian profile with  $7^\circ$  full width at half-maximum. Since the small-scale structure of the maps is dominated by uncorrelated instrument noise, we use a simple Gaussian approximation to the DMR beam instead of the Legendre polynomial expansion discussed in Wright et al. (1994). The pixels are searched in order of decreasing absolute temperature; we subtract each fitted Gaussian from the map before proceeding to the next pixel. To prevent contamination from Galactic plane sources,

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we only process pixels with Galactic latitude  $|b| > 30^\circ$ . Observation at three discrete frequencies provides further discrimination against high-latitude Galactic sources.

The fitted amplitudes  $T$  and uncertainties  $\delta T$  estimated from the known noise properties of the maps form a data set which we search for evidence of a population of faint unresolved sources. Most of the fitted amplitudes describe noise peaks or the short spatial frequency tail of the scale-invariant CMB temperature distribution. We test this hypothesis and search for any other source population by comparing the fitted amplitudes from the DMR maps to Monte Carlo simulations of instrument noise (top) or a superposition of noise and CMB anisotropy characterized by a Harrison-Zel'dovich spectrum with quadrupole normalization  $Q_{\text{rms-PS}} = 17 \mu\text{K}$ . We make a cut in the signal-to-noise ratio,  $T/\delta T$ , and bin the surviving amplitudes to form the integral distribution  $N(>T)$  of amplitudes brighter than threshold  $T$ . Different source populations are expected to have different distributions in temperature. Galactic emission (e.g., from compact H II regions) will appear as pixels of hotter than average temperature, while the Sunyaev-Zel'dovich effect creates a temperature decrement in this wavelength regime. We evaluate separately the integral distributions  $N(>T)$  for positive amplitudes and  $N(<T)$  for negative amplitudes, as well as the combined distribution  $N(>|T|)$  for the absolute value of the fitted amplitude.

Figures 1, 2, and 3 show the integral source distributions  $N(>T)$ ,  $N(<T)$ , and  $N(>|T|)$ , respectively, for the DMR 53 GHz data at signal to noise ratio  $T/\delta T > 1$ , compared to simulations of instrument noise (top) or a superposition of noise and CMB anisotropy (bottom). There are generally more "sources" detected in the  $(A+B)/2$  sum maps than the  $(A-B)/2$  difference maps, consistent with the expected increase

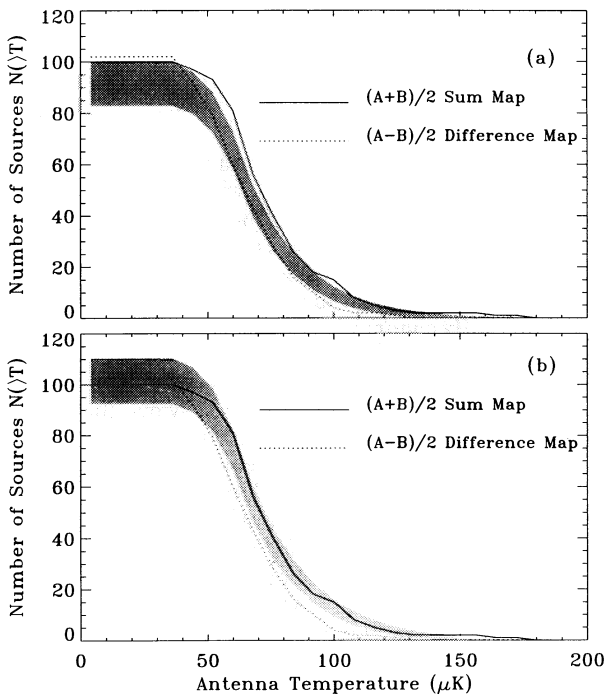


FIG. 1.—Integral source counts  $N(>T)$  for the DMR 53 GHz  $(A+B)/2$  sum maps and  $(A-B)/2$  difference maps at  $|b| > 30^\circ$  and signal-to-noise ratio  $T/\delta T > 1$ . The cut in signal-to-noise accounts for the flattening of the curves near zero amplitude. (a) Gray band is 68% confidence range of simulations of instrument noise. (b) Gray band is 68% confidence range of simulations of a superposition of instrument noise and scale-invariant CMB anisotropy normalized to  $Q_{\text{rms-PS}} = 17 \mu\text{K}$ .

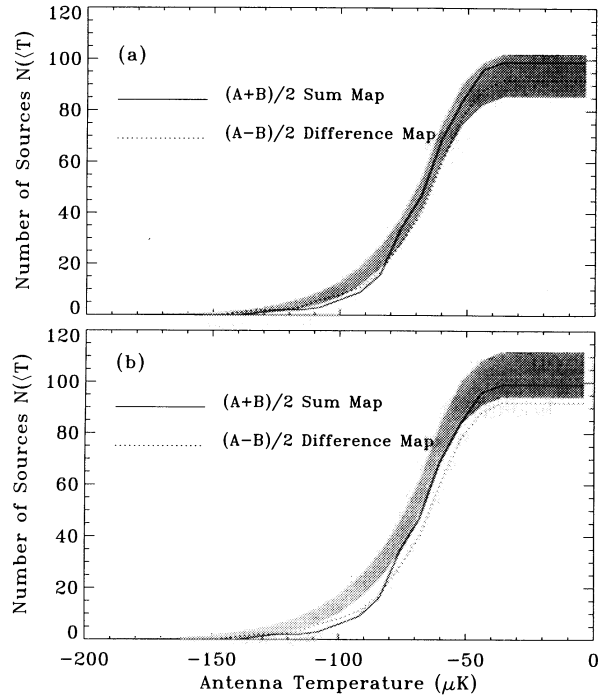


FIG. 2.—Integral source counts  $N(<T)$  at 53 GHz. Same parameters as Fig. 1.

in detections seen in the Monte Carlo simulations when CMB anisotropy is included. We quantify this comparison using the  $\chi^2$  statistic

$$\chi^2 = \sum_{i,j} (N_i - \langle N_i \rangle) (\mathbf{M}^{-1})_{ij} (N_j - \langle N_j \rangle),$$

where  $\langle N_i \rangle$  is the mean number of fitted sources in the  $i$ th temperature bin of the simulations and  $\mathbf{M}$  is the covariance

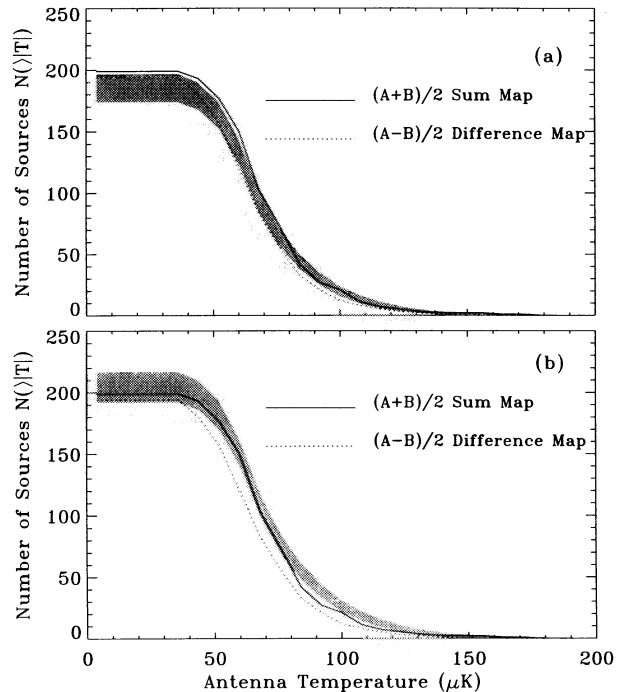


FIG. 3.—Integral source counts  $N(>|T|)$  at 53 GHz. Same parameters as Fig. 1.

TABLE 1  
COMPARISON OF FITTED AMPLITUDES TO SIMULATIONS<sup>a</sup>

DATA SET	MODEL = NOISE + CMB		MODEL = NOISE ONLY	
	(A + B)/2	(A - B)/2	(A + B)/2	(A - B)/2
31 GHz $N(> T)$ .....	14.7 (18%)	8.0 (73%)	15.7 (14%)	8.5 (66%)
31 GHz $N(< T)$ .....	10.3 (50%)	18.7 (7%)	10.5 (48%)	20.4 (4%)
31 GHz $N(>  T )$ .....	14.2 (29%)	21.0 (5%)	15.7 (21%)	22.8 (3%)
53 GHz $N(> T)$ .....	7.6 (83%)	23.7 (2%)	11.5 (38%)	22.0 (2%)
53 GHz $N(< T)$ .....	14.0 (29%)	7.8 (79%)	11.7 (37%)	4.3 (96%)
53 GHz $N(>  T )$ .....	11.3 (51%)	15.4 (20%)	12.1 (43%)	9.0 (70%)
90 GHz $N(> T)$ .....	17.0 (8%)	17.6 (7%)	24.9 (1%)	18.7 (5%)
90 GHz $N(< T)$ .....	3.1 (98%)	13.6 (19%)	2.4 (99%)	15.9 (10%)
90 GHz $N(>  T )$ .....	5.3 (91%)	14.5 (19%)	9.2 (61%)	19.7 (4%)

<sup>a</sup>  $\chi^2$  statistic and probability for a random realization to have larger  $\chi^2$  than the DMR value. Fitted amplitudes from the DMR maps are compared to Monte Carlo realizations for  $|b| > 30^\circ$  and signal-to-noise ratio  $T/\delta T > 1$ .

matrix between bins. We neglect the high-temperature tail of the distribution ( $\langle N_i \rangle < 5$ ) where Poisson statistics dominate, and use an independent test for rare bright sources.

Table 1 shows the  $\chi^2$  statistic of the DMR sum and difference maps at signal-to-noise ratio  $T/\delta T > 1$ , compared to simulations with noise or a superposition of noise and scale-invariant CMB anisotropy. The probability of obtaining a random realization with  $\chi^2$  larger than the DMR value is also shown. We repeat the analysis for several thresholds in signal-to-noise ratio with generally consistent results. The sum maps have somewhat better  $\chi^2$  when compared to simulations which include CMB anisotropy than to a blank sky, but the difference is small, reflecting the low DMR response to multipoles  $l > 20$ . The apparent decrement of negative sources (Fig. 2b) is not statistically significant. There is no detection of a population of unresolved sources in the DMR data.

Table 2 compares the total number of positive and negative fitted source amplitudes for the sum and difference maps compared to simulations. There is no evidence for any statistically significant asymmetry between positive and negative amplitudes, in agreement with the lack of statistically significant skewness in the raw DMR pixel temperatures (Smoot et al. 1994). Unresolved Galactic or extragalactic sources, both of which are asymmetric in temperature, contribute negligibly to the anisotropy on  $7^\circ$  scales observed by DMR.

The high-temperature tail of the integral source distribution provides an upper limit to the brightest potential sources in the DMR data. Table 3 shows the brightest fitted source in the sum and difference maps, along with the temperature for which the mean number of sources in the simulated data, plus twice the standard deviation, equals unity. Sources brighter than this limit, should any exist, would be detected. The brightest pixels are in different positions in each map and are consistent with the noise properties of the data. There are no unresolved sources at limit  $\Delta T < 179 \mu\text{K}$  (95% CL) antenna temperature in the most sensitive 53 GHz maps.

### 3. COMPARISON TO KNOWN SOURCES

Bennett et al. (1993) cross-correlate the first-year DMR maps with maps of known noncosmological sources to place an upper limit on the total contribution of these sources to the anisotropy in the DMR data. We consider here the related question of whether the brightest of these sources are detected in the DMR data, even if the spatial average over all such sources is not.

Inverse Compton scattering of CMB photons from hot intracluster gas causes a decrement in the Rayleigh-Jeans part of the CMB spectrum,  $\Delta T = -2yT_0$ , where  $T_0$  is the unperturbed CMB temperature and the Comptonization parameter  $y$  is proportional to the integrated electron pressure along the

TABLE 2  
POSITIVE AND NEGATIVE SOURCE COUNTS IN DMR MAPS

Data Set	$N_+$ <sup>a</sup>	$N_-$ <sup>a</sup>	$\langle T_+ \rangle^b$ ( $\mu\text{K}$ )	$\langle T_- \rangle^b$ ( $\mu\text{K}$ )
31 GHz (A + B)/2 .....	98	109	202	-197
31 GHz (A - B)/2 .....	90	107	199	-194
Model (Noise + CMB) .....	93 $\pm$ 8	94 $\pm$ 8	201 $\pm$ 7	-202 $\pm$ 7
Model (Noise only) .....	90 $\pm$ 8	94 $\pm$ 8	202 $\pm$ 7	-202 $\pm$ 7
53 GHz (A + B)/2 .....	100	99	73	-65
53 GHz (A - B)/2 .....	102	92	62	-65
Model (Noise + CMB) .....	101 $\pm$ 8	103 $\pm$ 9	70 $\pm$ 2	-70 $\pm$ 2
Model (Noise only) .....	91 $\pm$ 8	93 $\pm$ 8	69 $\pm$ 2	-68 $\pm$ 2
90 GHz (A + B)/2 .....	114	86	101	-95
90 GHz (A - B)/2 .....	86	102	100	-94
Model (Noise + CMB) .....	96 $\pm$ 8	95 $\pm$ 8	97 $\pm$ 3	-96 $\pm$ 3
Model (Noise only) .....	93 $\pm$ 8	90 $\pm$ 8	97 $\pm$ 3	-95 $\pm$ 3

<sup>a</sup> Total number of sources at  $|b| > 30^\circ$  and signal-to-noise ratio  $T/\delta T > 1$ .

<sup>b</sup> Mean antenna temperature of fitted source amplitudes.

TABLE 3  
BRIGHTEST FITTED SOURCES IN DMR MAPS<sup>a</sup>

Frequency (GHz)	Sum Map ( $\mu\text{K}$ )	Difference Map ( $\mu\text{K}$ )	CMB + Noise Simulation ( $\mu\text{K}$ )	Noise Simulation ( $\mu\text{K}$ )
31.....	531	502	527	542
53.....	171	-138	179	171
90.....	206	-266	241	236

<sup>a</sup> Fitted amplitudes at  $|b| > 30^\circ$  and signal-to-noise ratio  $T/\delta T > 1$ . Amplitudes are antenna temperature.

line of sight (Sunyaev & Zel'dovich 1980). The Sunyaev-Zel'dovich (S-Z) effect has been observed toward a number of X-ray-selected clusters (Jones et al. 1993; Grainge et al. 1993; Herbig et al. 1994; see Birkinshaw 1990 for a recent review) with typical amplitude  $\Delta T \approx -400 \mu\text{K}$  on arcminute angular scales. Although beam dilution makes these unlikely to be observed in the DMR data, we have examined the DMR maps for evidence of isolated S-Z sources. Table 4 lists the amplitudes fitted to the DMR sum and difference maps at the positions of firm S-Z detections on arcminute angular scales (0016+16, A665, A773, and A2218). There is no statistically significant decrement toward these sources.

Nearby clusters are more likely to contribute a detectable signal to the DMR maps given their larger angular extent. We have searched the DMR maps toward nearby clusters at  $|b| > 20^\circ$  (Coma, Virgo, Perseus-Pisces, Hercules, and Hydra) for evidence of Sunyaev-Zel'dovich cooling indicative of an extended low surface density cluster atmosphere. To account for diffuse low surface density Comptonization (which may be ill described by a Gaussian profile), results for these clusters are from a "ring" technique, in which the temperature of all pixels within a disk of radius  $\theta$  is compared to the annulus of equal area surrounding the disk. There is no detection of the Sunyaev-Zel'dovich cooling as  $\theta$  is varied from  $5^\circ$  to  $20^\circ$ . Table 4 lists the results for  $\theta = 10^\circ$ , placing a limit  $\delta y < 7.3 \times 10^{-6}$  (95% CL) for the Comptonization from nearby clusters. Rephaeli (1993) has recently analyzed the *HEAO 1*-A2 database to search for evidence of gas associated with superclusters of gal-

axies. He finds no evidence for such gas, and uses the upper limit on such supercluster gas to infer an upper limit on the S-Z contribution to the DMR maps from such sources of  $\delta y > 10^{-7}$ . This limit is in agreement with the upper limit  $\delta y < 7.3 \times 10^{-6}$  (95% CL) imposed by direct examination of the DMR data toward nearby clusters (Table 4), or the limit  $\delta y < 2 \times 10^{-6}$  (95% CL) from the full cross-correlation of the DMR maps with either the rich cluster distribution or the *HEAO 1*-A2 data (Bennett et al. 1993).

Extragalactic radio sources are expected to contribute negligibly to the DMR signal (Franceschini et al. 1989). Table 5 shows the fitted amplitudes toward the positions of the largest apparent diameter objects (the Large Magellanic Cloud and M31) and the brightest sources at 53 GHz from a compilation of radio surveys and millimeter catalogs with spectra interpolated to the DMR wavelengths (Kühr et al. 1981; Herbig & Readhead 1992; Edelson 1987; Steppe et al. 1992, 1988). The LMC is expected to be the brightest extragalactic source, contributing approximately  $50 \mu\text{K}$  in a single pixel at 53 GHz (Bennett et al. 1993). The  $2\sigma$  upper limit of  $87 \mu\text{K}$  from Table 5 is consistent with this level of emission. As expected, fitting individual pixels shows no detection of unresolved extragalactic radio sources in the DMR data.

#### 4. DISCUSSION

There is no evidence in the DMR data for unresolved sources beyond the contribution expected from the high spatial frequency tail of a scale-invariant spectrum of anisotropy. The distribution of amplitudes fitted to a Gaussian profile at each pixel is symmetric about zero, implying that unresolved Galactic sources contribute negligibly compared to the CMB and noise in the two-year maps. The high-temperature tail of the fitted amplitude distribution for  $|b| > 30^\circ$  is consistent with the noise properties of the data; there are no sources brighter than  $179 \mu\text{K}$  antenna temperature at 53 GHz (95% CL), corresponding to a flux density  $322 \text{ Jy}$ .

TABLE 4

FITTED AMPLITUDES TOWARD SUNYAEV-ZEL'DOVICH CANDIDATES ( $\mu\text{K}$ )

Source	$T_{31}$	$T_{53}$	$T_{90}$	Weighted mean
0016+16 <sup>a</sup> .....	-40 ± 182	-7 ± 60	-61 ± 98	-23 ± 49
	-46 ± 182	-12 ± 60	-47 ± 98	-23 ± 49
A665 <sup>a</sup> .....	-29 ± 120	12 ± 48	-5 ± 78	3 ± 38
	68 ± 120	34 ± 48	-8 ± 78	27 ± 38
A773 <sup>a</sup> .....	31 ± 138	-13 ± 52	-58 ± 85	-20 ± 42
	-13 ± 138	24 ± 52	0 ± 85	15 ± 42
A2218 <sup>a</sup> .....	-179 ± 115	12 ± 45	28 ± 74	-3 ± 36
	-38 ± 115	-12 ± 45	-10 ± 74	-14 ± 36
Coma <sup>b</sup> .....	36 ± 67	10 ± 23	-19 ± 37	5 ± 19
	-102 ± 67	-14 ± 23	16 ± 37	-13 ± 19
Virgo <sup>b</sup> .....	26 ± 76	-48 ± 24	26 ± 39	-24 ± 20
	60 ± 76	13 ± 24	-44 ± 39	2 ± 20
Perseus-Pisces <sup>b</sup> ...	112 ± 87	29 ± 34	27 ± 54	37 ± 27
	-2 ± 87	3 ± 34	-15 ± 54	-2 ± 27
Hercules <sup>b</sup> .....	39 ± 59	-15 ± 25	-11 ± 40	-8 ± 20
	-117 ± 59	11 ± 25	59 ± 40	8 ± 20
Hydra <sup>b</sup> .....	-137 ± 83	-21 ± 29	-19 ± 47	-30 ± 24
	12 ± 83	-5 ± 29	-16 ± 47	-6 ± 24

<sup>a</sup> First line for each source refers to the fitted amplitudes in the  $(A+B)/2$  sum map, while the second line refers to the  $(A-B)/2$  difference map. All amplitudes are thermodynamic temperature.

<sup>b</sup> Thermodynamic temperature differences from "ring" technique.

TABLE 5

FITTED AMPLITUDES TOWARD RADIO SOURCES ( $\mu\text{K}$ )

Source	$T_{31}$	$T_{53}$	$T_{90}$
LMC <sup>a</sup> .....	38 ± 106	11 ± 38	0 ± 55
	125 ± 106	3 ± 38	31 ± 55
M31 .....	92 ± 162	61 ± 57	-7 ± 82
	63 ± 162	2 ± 57	-80 ± 82
3C 273.0 .....	0 ± 189	11 ± 53	-8 ± 77
	-30 ± 189	21 ± 53	-19 ± 77
3C 345 .....	-42 ± 100	33 ± 37	-6 ± 54
	-19 ± 100	-27 ± 37	8 ± 54
0537-441 .....	12 ± 100	0 ± 33	-1 ± 47
	71 ± 100	6 ± 33	39 ± 47

<sup>a</sup> First line for each source refers to the fitted amplitudes in the  $(A+B)/2$  sum map, while the second line refers to the  $(A-B)/2$  difference map. All amplitudes are antenna temperature.

A search of pixels where extragalactic radio sources are expected to be brightest shows no evidence of emission, nor was any expected. A search toward bright nearby clusters shows no evidence for Sunyaev-Zel'dovich cooling and places a limit on the Comptonization parameter  $\delta y < 7.3 \times 10^{-6}$  (95% CL) for these clusters. Since the Sunyaev-Zel'dovich effect is independent of the cluster redshift (except for the angular size), clusters undetectable in the optical or X-ray at high redshifts could in principle affect the DMR maps. The absence of any significant S-Z effect in the DMR maps (e.g., the lack of significant temperature asymmetry) is evidence against significant contamination of the DMR maps from high-redshift clusters.

The recent report of strong unresolved sources at  $0.5^\circ$  angular scale has kindled interest in non-Gaussian models of CMB anisotropy. The Medium Scale Anisotropy Measurement (MSAM) has detected two unresolved sources (width in right ascension  $< 15'$ ) consistent with CMB features of amplitude 4 Jy at 1.8 mm wavelength, or thermodynamic temperature  $\Delta T \approx 250 \mu\text{K}$  (Cheng et al. 1994). These features, if real, may imply a non-Gaussian component to the CMB anisotropy on  $0.5^\circ$  angular scale. Because of beam dilution, the lack of unresolved Planckian sources in the DMR data cannot provide meaningful limits to measurements on smaller angular scales. The MSAM sources would contribute less than  $2 \mu\text{K}$  to the DMR maps, well below the  $90 \mu\text{K}$  mean noise level per pixel of the DMR 53 GHz sum maps. Conversely, an unre-

solved source at the DMR upper limit of 322 Jy would dominate a survey at  $0.5^\circ$  angular resolution and so must be rare to avoid detection in surveys of limited sky coverage. If a population of Planck spectrum sources with angular size less than  $0.5^\circ$  were to exist, the integral number density per unit area brighter than flux density  $S$  assumed to vary as a power law  $n(>S) \sim S^p$ , we may normalize the assumed distribution to the MSAM source density (two sources in 0.002 sr at 4 Jy) and derive an upper limit for larger  $S$  based on the DMR results (less than one source in  $2\pi$  sr at 322 Jy) with the result  $n(>S) < 2 \times 10^4 (S/1 \text{ Jy})^{-2} \text{ sr}^{-1}$ . Interpretation of the MSAM measurement or other recent degree-scale anisotropy results (ACME-HEMT, Schuster et al. 1993; MAX, Gundersen et al. 1993; Big Plate, Wollack et al. 1993) in terms of a population of unresolved sources is complicated by the small sampling of the sky (typically 0.1%) and the sparse spatial sampling of the instrumental scan patterns. A CMB survey with  $0.5^\circ$  resolution or better covering a significant fraction of the sky would be immensely valuable regarding the question of whether such sources exist.

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